Figure from Kinematics and Evolution of the Giant HH34 Complex
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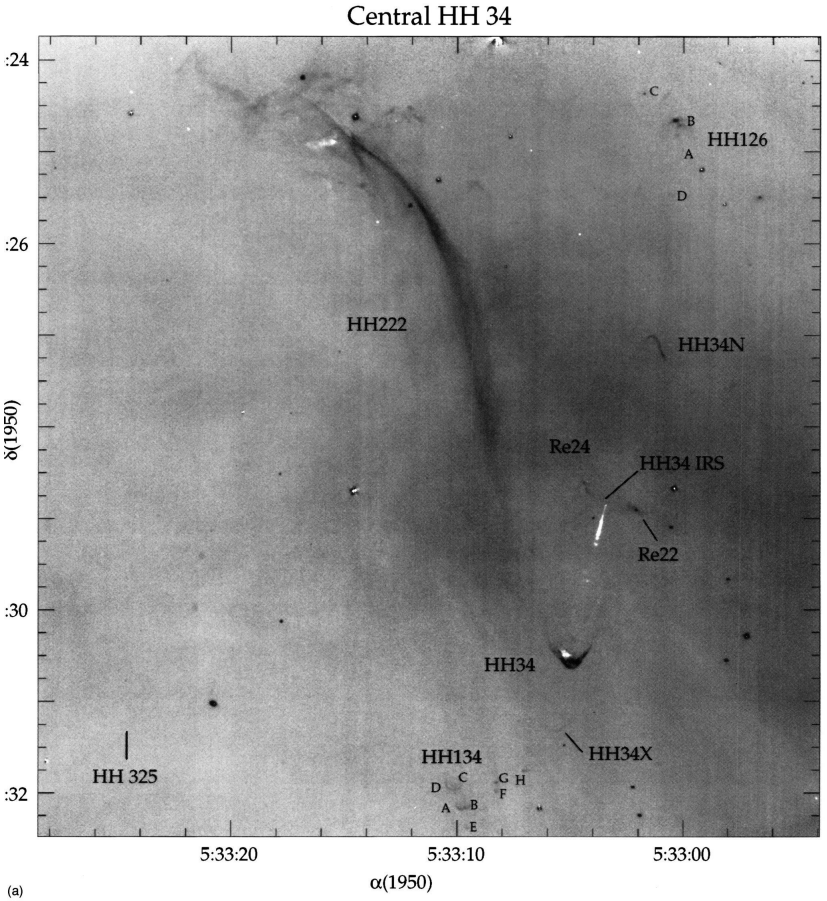

3. (a) The difference (Hα−[S II]) images for the individual NTT pointings in the HH34 complex. Hα is dark and [S II] is bright. (b) The region south of HH34 IRS. (c) The region north of HH34 IRS.

3, and in the HL/XZ Tau region in the L1551 cloud. wever, there are arguments for including HH85, HH40, HH33 in the HH34 flow, as outlined below.

HH85: The HH85 complex can be paired with the HH173 plex to the south. Both complexes lie at roughly the e distance from HH34 IRS, and have similar radial ve- ty magnitudes relative to the HH34 IRS core. They also e similar morphologies, although HH85 is brighter than regular spacing between HH126, HH85E, HH85B, and HH40 suggests a common source. Since HH34 IRS is almost certainly the source of HH126, it is likely to be the source of the other northern HH objects as well.

HH33/40: The emission connecting HH40 to HH33 has long been believed to trace a jet (Mundt *et al.* 1984, 1987; however, see Eislöffel & Mundt 1997). This is consistent with our measured proper motions, which are directed